

# Towards a Newtonian softening length in discs simulations?

Audrey Trova<sup>1</sup>, Jean-Marc Huré<sup>2,3</sup>, Vladimir Karas<sup>1</sup>

<sup>1</sup> Astronomical Institute of the Academy of Sciences, Boční II 1401, CZ-14100 Prague, Czech Republic

<sup>2</sup> Univ. Bordeaux, LAB, UMR 5804, F-33270 Floirac, France,

<sup>3</sup> CNRS, LAB, UMR 5804, F-33270, Floirac, France

## Abstract

In numerical simulations of discs, self-gravitating potentials and forces are traditionally computed from softened gravity (see **panel 1**). In this context, the softening length  $\lambda$  is set to a fraction of the disc local thickness  $h$ , typically  $\lambda/h \in [0.3 - 1.2]$ . As shown (see **panel 2**), such a prescription is not appropriate. We have computed the gravitational potential of a numerical cell, and deduced the corresponding  $\lambda$ -value. It turns out that the length not only depends on the shape of the cell but it can be an imaginary number (see **panels 3 and 4**). A dipolar expansion shows that  $\lambda$  effectively does not depend on the cell's height only (see **panel 5**). We present a novel prescription, valid at long-range, that preserves the Newtonian properties at the scale of the numerical grid cells. A general analysis is in progress.

## 1. The softened point mass potential

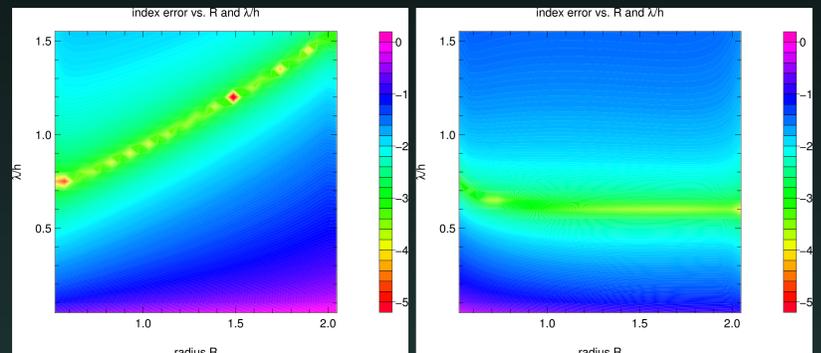
The softened potential is used in continuous media to simplify the numerical treatment of Newton's triple integral. It avoids the kernel singularity through a modification of the relative separation, namely

$$|\vec{r} - \vec{r}'| \rightarrow \sqrt{|\vec{r} - \vec{r}'|^2 + \lambda^2}, \quad (1)$$

where  $\lambda$  is called the "softening length". This free parameter is selected by authors more or less arbitrarily [Papaloizou and Lin, 1989, Morishima and Saio, 1994, Baruteau and Masset, 2008, Meru and Bate, 2012]. In a discretized disc, the potential of each cell is replaced by the softened point mass potential (also known as Plummer potential), namely

$$\psi_{\text{Plum.}}(\vec{r}; \vec{r}_0, \lambda) = -\frac{Gm}{\sqrt{|\vec{r} - \vec{r}_0|^2 + \lambda^2}}. \quad (2)$$

The value of  $\lambda$  has a severe impact on the simulations. The figure 1 shows the error on the gravitational potential as function of the radius and ratio  $\lambda/h$  in two concrete cases. We see that the nominal value of  $\lambda$  depends on the radius in the disc, and on density profile as well, which is not acceptable.



**Figure 1:** Relative error (log. scale) on potential values when computing the potential in a flat homogenous disc (*left*) and in a flared power-law disc (*right*) when using the softened potential with  $\lambda/h = cst$ . Parameters and setup: disc inner edge at **0.5**; discretization into **32 × 64** cells in the  $(R, \theta)$ -plane; regular spacing;  $h \propto a$  and mass density  $\rho \propto a^{-2.5}$  for the flared power-law disc.

## 2. $\lambda$ can be an imaginary number !

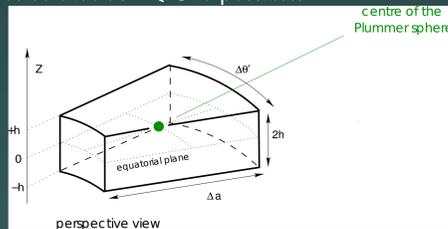
According to the Newtonian theory, the potential of a cylindrical sector, as depicted in figure 2, is given by the triple integral

$$\psi_{\text{cell}}(\vec{r}) = -\frac{Gm}{V} \int_{z_0-h}^{z_0+h} dz \int_{a_0-\frac{1}{2}\Delta a}^{a_0+\frac{1}{2}\Delta a} da \int_{\theta_0-\frac{1}{2}\Delta\theta}^{\theta_0+\frac{1}{2}\Delta\theta} ad\theta' \frac{1}{|\vec{r} - \vec{r}'|}. \quad (3)$$

So,  $\lambda$  reproduces exactly the Newtonian potential of the cell if

$$\psi_{\text{Plum.}}(\vec{r}; \lambda) - \psi_{\text{cell}}(\vec{r}) = 0, \quad \text{i.e.} \quad \lambda^2 = \left(\frac{Gm}{\psi_{\text{cell}}}\right)^2 - |\vec{r} - \vec{r}_0|^2, \quad (4)$$

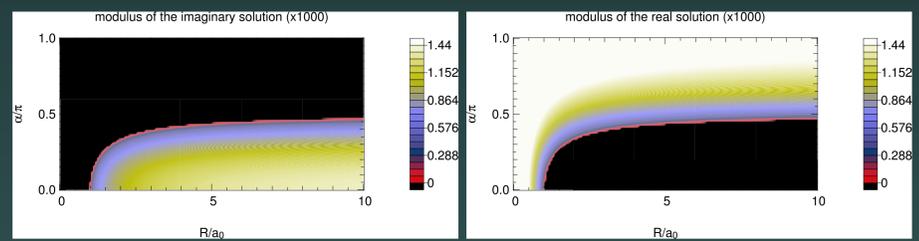
everywhere in space. We notice that  $\lambda^2 < 0$  is possible.



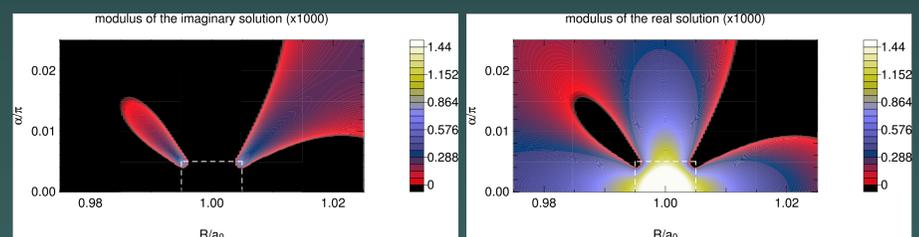
**Figure 2:** Cylindrical cell (3D-view and projection in the midplane), centre of the Plummer sphere and associated notations. Points A to E mark the centre of neighbouring cells.

We have computed  $\lambda$  from Eq.(4) using the contour integral given by [Huré et al., 2014], see figure on the right.

It turns out that  $\lambda$  can be an **imaginary number** !



**Figure 3:** Value of  $\lambda$  computed in the equatorial plane from Eq.(4): imaginary solutions (*left*) and real solutions (*right*). The coordinates of the cell's centre are  $R/a_0 = 1$ ,  $\alpha = 0$  and  $z_0$ .



**Figure 4:** Same legend and same color code as for Fig. 3, but zoomed around the **numerical cell** (boundary in dashed white line).

## 3. $\lambda$ at long-range from a dipolar expansion

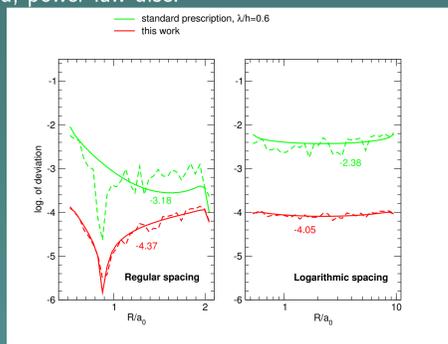
We can expand the kernel  $1/|\vec{r} - \vec{r}'|$  over  $\alpha = \theta - \theta_0$  before performing the integrations in Eq.(3). We then focus on the long-range behavior of  $\psi_{\text{cell}}$  in Eq.(4). After some tedious calculus, we obtain the following approximation for  $\lambda$ :

$$\frac{\lambda^2}{4a_0R} \approx \frac{\Delta\theta'^2}{48} \cos \alpha - \frac{1}{24} \left(\frac{\Delta a}{a_0}\right)^2 \left[ \cos \alpha - \frac{3a_0(\zeta_0^2 + R^2 \sin^2 \alpha)}{2|\vec{r} - \vec{r}_0|^2} \right] + \frac{h^2}{12a_0R} \equiv \bar{\lambda}^2 \quad (5)$$

We conclude that  $\lambda$  is not proportional to  $h$ , as often considered.

## Conclusion & perspectives

With such an improved prescription, the error on potential values is reduced by **2 – 3** orders of magnitude typically with respect to the standard prescription where  $\lambda \propto h$ . This is illustrated in the figure below for a flared, power-law disc.



To conclude, we show that

- the standard prescription  $\lambda \propto h$  is not the nominal choice.
- the appropriate softening length can be an imaginary number; this corresponds to a point mass potential weaker than that of the cylindrical cell.
- $\lambda$  is a complicated function of space and cell's geometry. Collective effects show  $\lambda$  depends on how the disc is discretized (radial, azimuthal and vertical sampling).
- it is possible to determine an approximation for the softening length valid at long-range. Thus, this formula enables to mimic the Newtonian potential of a cylindrical from the Plummer potential, at long-range.

More details are given in Huré and Trova [2014].

This work can be continued in several ways. It could be interesting to produce a complete analytical value of  $\lambda$  (valid at short-range too), which is tricky, and to see how this new prescription impacts on hydrodynamical simulations.

## References

- C. Baruteau and F. Masset. Type I Planetary Migration in a Self-Gravitating Disk. *ApJ*, 678:483–497, May 2008. doi: 10.1086/529487.
- J.-M. Huré and A. Trova. A truly newtonian softening length in disc simulations. Manuscript submitted for publication to MNRAS, 2014.
- J.-M. Huré, A. Trova, and F. Hersant. Self-gravity in curved mesh elements. *Celestial Mechanics and Dynamical Astronomy*, February 2014. doi: 10.1007/s10569-014-9535-x.
- F. Meru and M. R. Bate. On the convergence of the critical cooling time-scale for the fragmentation of self-gravitating discs. *MNRAS*, 427:2022–2046, December 2012. doi: 10.1111/j.1365-2966.2012.22035.x.
- T. Morishima and H. Saio. Viscous Evolution of Self-Gravitating Galactic Disks Under Modified Newtonian Dynamics. *MNRAS*, 267:766–+, April 1994.
- J. C. B. Papaloizou and D. N. C. Lin. Nonaxisymmetric instabilities in thin self-gravitating rings and disks. *ApJ*, 344:645–668, September 1989. doi: 10.1086/167832.